The Search for Life Finding Earth 2.0 Dr Jonathan Crass



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Are we alone?

Searching for Exoplanets

Searching for Exoplanets

There are four major techniques we use to search for exoplanets:

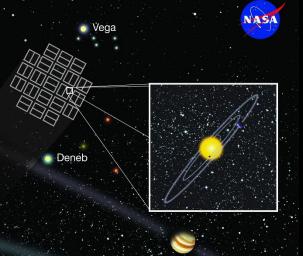
- Transit Technique
- Radial Velocity Technique
- Gravitational Microlensing
- Direct Imaging



National Aeronautics and Space Administration

Kepler

NASA's First Mission Capable of Finding Earth-size & Smaller Planets

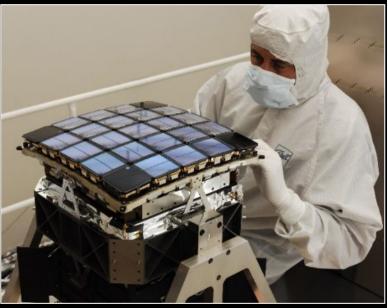


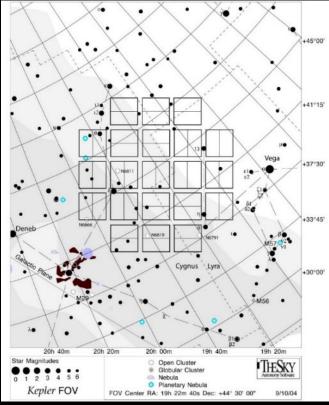
WARNING: OBJECTS IN THIS RENDITION APPEAR LARGER AND CLOSER TOGETHER THAN THEY ARE IN REALITY.

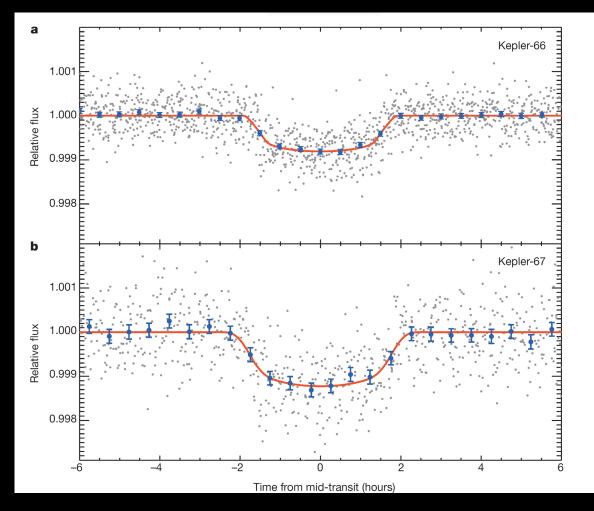


The Kepler Mission

R

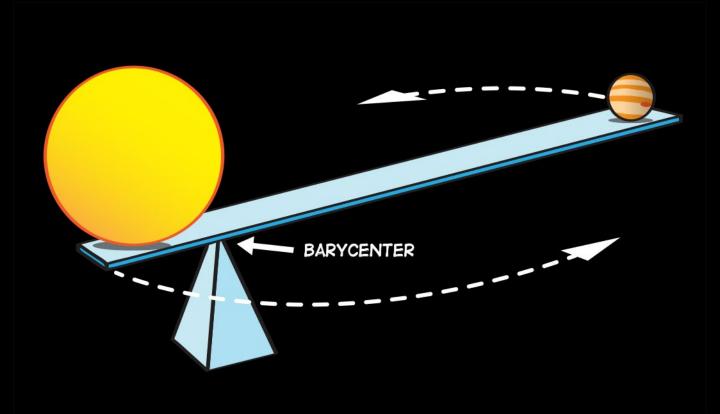


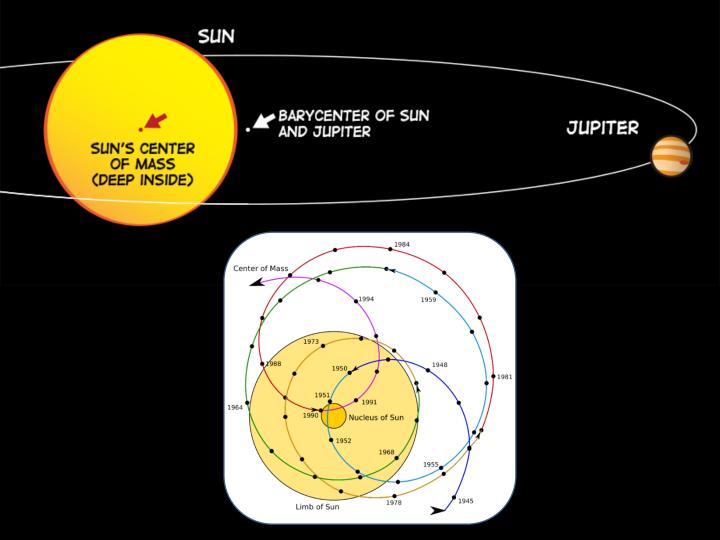


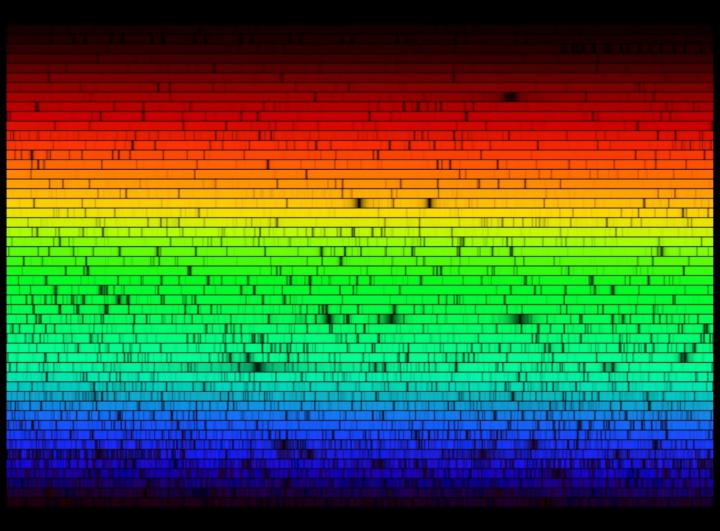




TRANSITING EXOPLANET SURVEY SATELLITE DISCOVERING NEW EARTHS AND SUPER-EARTHS IN THE SOLAR NEIGHBORHOOD

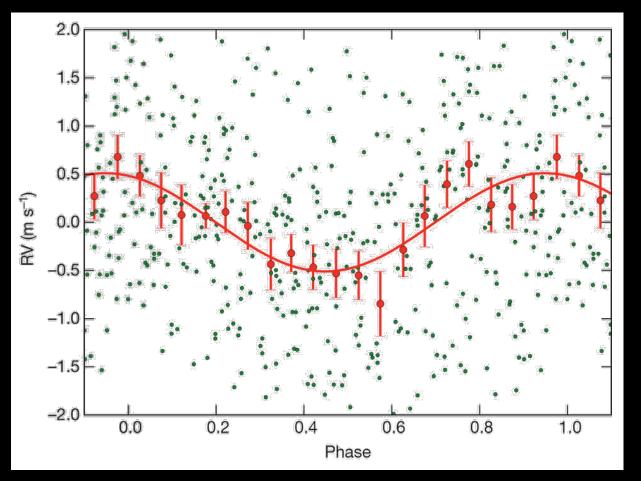




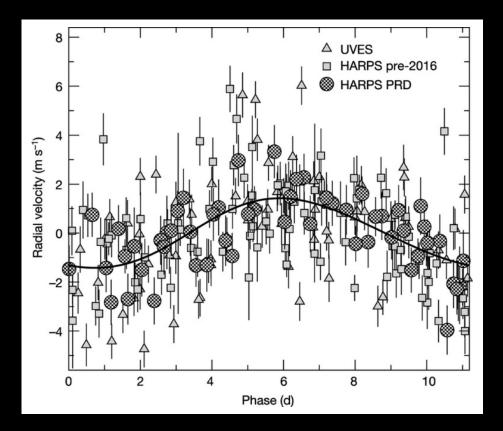








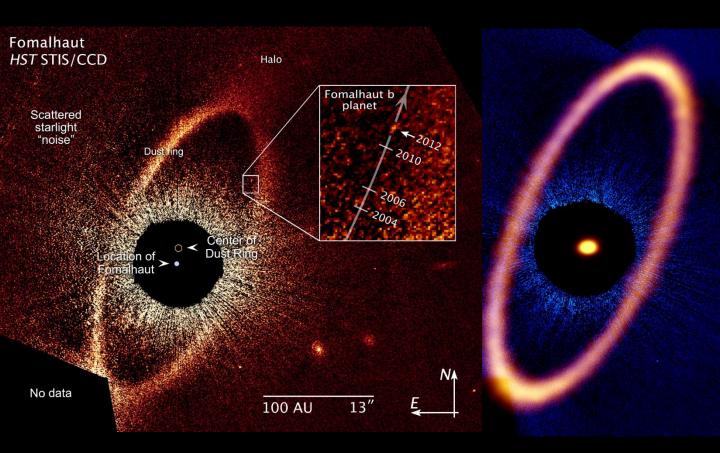
Proxima Centauri b

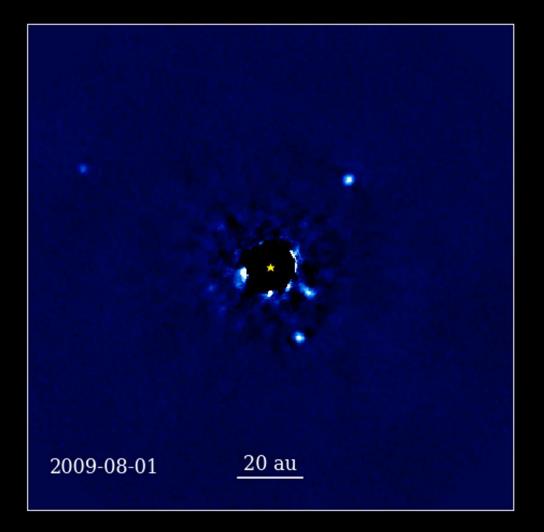












A very new field...

- 1988 A possible planetary detection...
- 1992 The discovery of a planet around a Pulsar
- 1995 The discovery of a planet around a Sun-like star: 51 Pegasi
- 2016 Earth-Size Planet in Habitable Zone: Proxima Centauri
- 2017 7 Earth-Size Planets around a single star: TRAPPIST-1

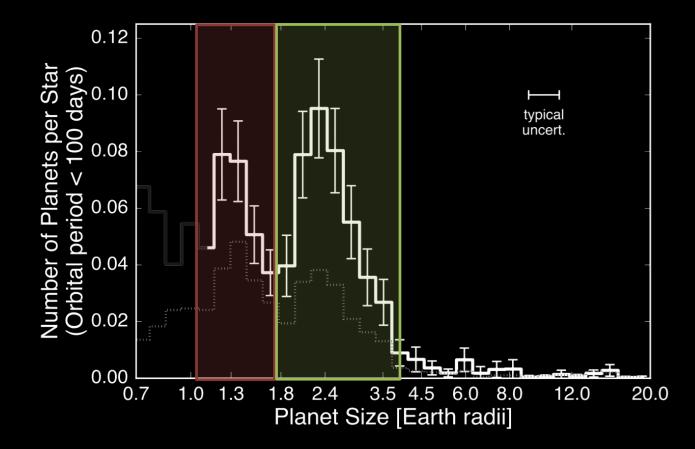


James Peebles Michel Mayor Didier Queloz

5811 CONFIRMED EXOPLANETS

Detection Method	Number of planets	Detection Method	Number of planets
Astrometry	3	Imaging	82
Radial Velocity	1096	Transit	4329
Transit Timing	32	Eclipse Timing	17
Microlensing	232	Pulsar Timing	8
Orbital Brightness Modulation	9	Other	3

Data from NASA Exoplanet Archive. Image credits: ESO & NASA







The signatures of life

ABSORPTION AND EMISSION SPECTRA

550 nm 650 nm Absorption NITROGEN Emission Absorption OXYGEN Emission

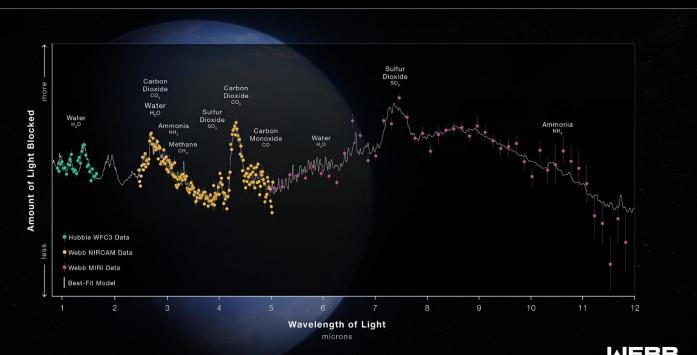
Water Oxygen Carbon Dioxide Methane



GAS-GIANT EXOPLANET WASP-107 b TRANSMISSION SPECTRUM

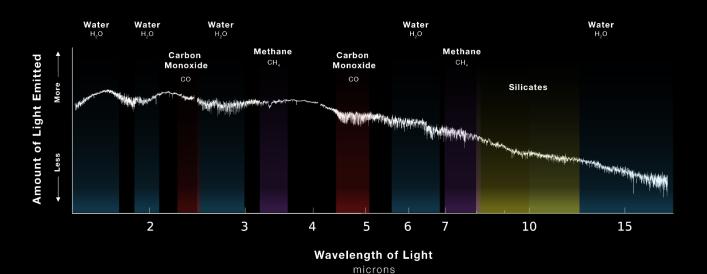
Hubble WFC3 | Grism Spectroscopy Webb NIRCam | Grism Spectroscopy Webb MIRI | Low-Resolution Spectroscopy

SPACE TELESCOPE

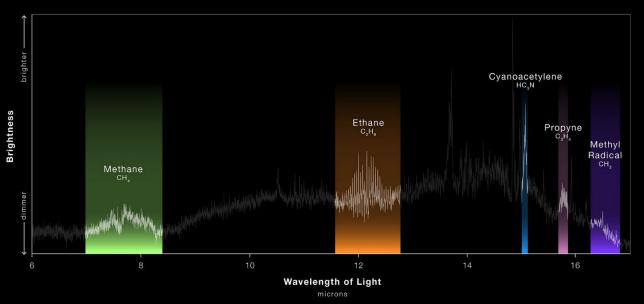


EMISSION SPECTRUM

NIRSpec and MIRI | IFU Medium-Resolution Spectroscopy



VERY LOW-MASS STAR : ISO-CHAI 147 HYDROCARBONS IN PROTOPLANETARY DISK





MIRI | Medium Resolution Spectroscopy

Life in our Galaxy

The Drake Equation

The Drake Equation

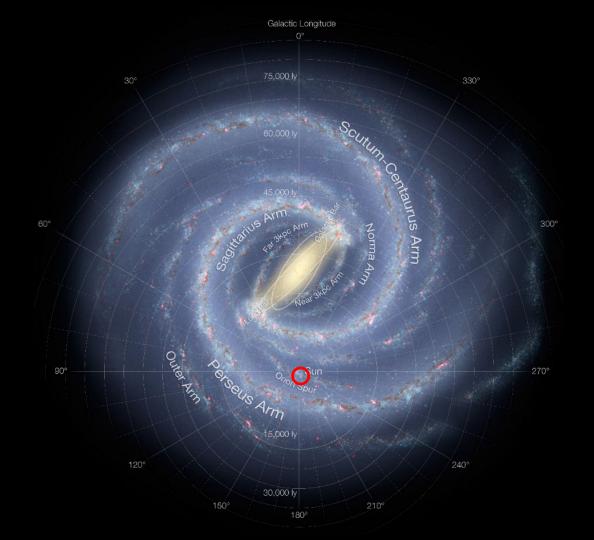
$$N = R^* \times f_p \times n_e \times f_l \times f_i \times f_c \times L$$

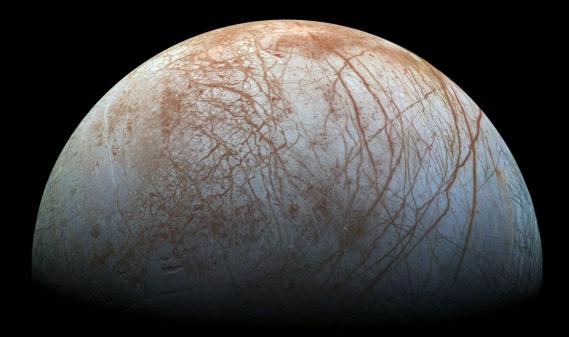
N = The number of civilizations in The Milky Way Galaxy R* = The rate of formation of stars f_p = The fraction of those stars with planetary systems n_e = The number of planets, per solar system, which can support life f_1 = The fraction of suitable planets on which life actually appears f_i = The fraction of life bearing planets on which intelligent life emerges f_c = The fraction of civilizations that develop technology to emit signals into space

L = The length of time such civilizations release these signals into space



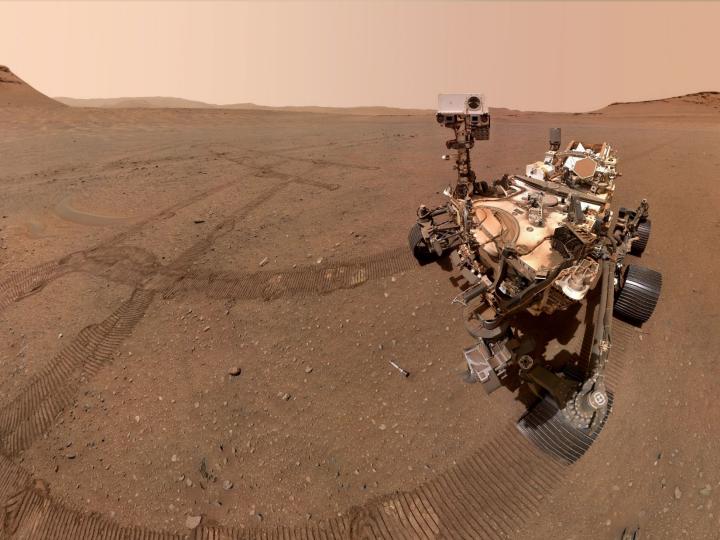






Life in our own Solar System



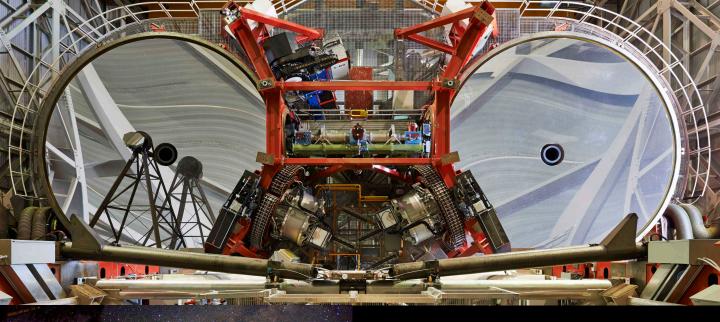




The Future

The remaining questions

Is our Solar System special? How do exoplanets form? What are exoplanets made of? Do they show signatures of life?



iLocater





The National Academies of SCIENCES • ENGINEERING • MEDICINE

Pathways to Discovery in Astronomy and Astrophysics for the 2020s



Theme #1 - Worlds and Suns in Context

"Pathways to Habitable Worlds" is a step-by-step program to identify and characterize Earth-like extrasolar planets, with the ultimate goal of obtaining imaging and spectroscopy of potentially habitable worlds.



